

# Cosmological Gravitational Waves (CosmoGW) as a probe of the early universe

High Energy Physics group at the Institute for Theoretical Physics at  
TU Wien (Feb. 20, 2026)



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SNSF Ambizione grant (2023–2026): “*Exploring the early universe with gravitational waves and primordial magnetic fields.*”

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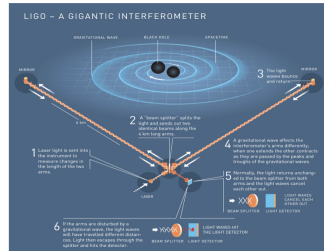
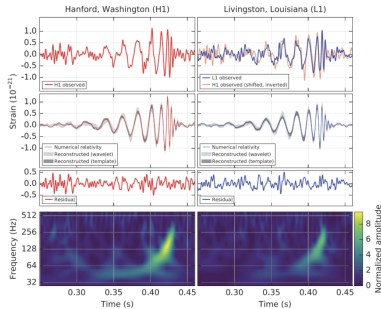
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Professor at University of Geneva  
and staff member at CERN



# First detections of gravitational waves

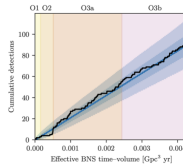
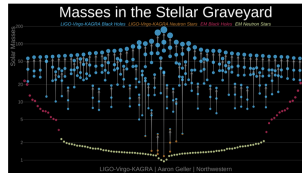
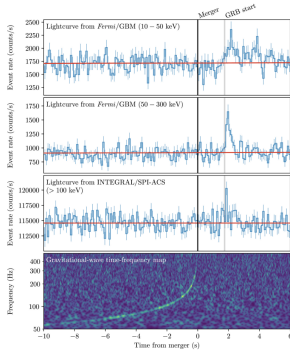
- Gravitational waves are opening a new window into our understanding of the Universe
  - First event GW150914 detected by LIGO-Virgo collaboration<sup>1</sup>



<sup>1</sup>[LIGO-Virgo Collaboration], *Phys. Rev. Lett.* **116**, 061102 (2016)

# Introduction and Motivation

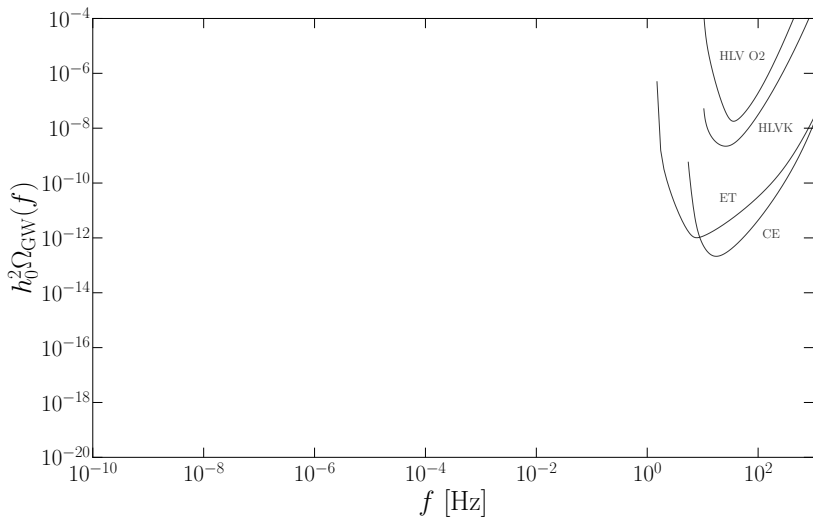
- GW170817 NS binary merger<sup>2</sup> first detection of GW and EM counterpart (constraint on the GW speed, measure of the Hubble rate, neutron star equation of state, ...)
- Several following events: LIGO-Virgo-KAGRA concluded the fourth observing run in November 2025 →  $\sim 340$  events up to O4 (90 up to O3b)<sup>3</sup> and 250 during O4.



<sup>2</sup>[LIGO-Virgo-Fermi GBM-Integral collaborations], Astrophys.J.Lett. 848 (2017) 2, L13

<sup>3</sup>[LIGO-Virgo Collaboration], GWTC-3, arXiv:2111.03606 (2021).

## Gravitational spectrum (ground-based detectors)





# LISA

- Laser Interferometer Space Antenna (LISA) is a space-based GW detector
- Approved in 2017 as one of the main research missions of ESA (L3) with NASA collaboration
- Launch planned for 2035
- Composed by three spacecrafts in a distance of 2.5M km
- LISA cosmology working group (since 2015,  $\sim 230$  members)

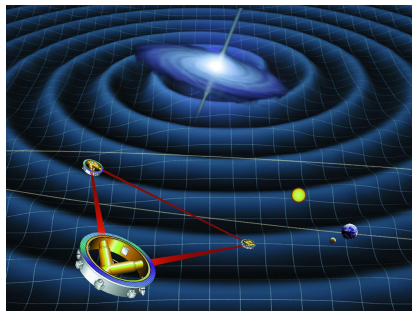
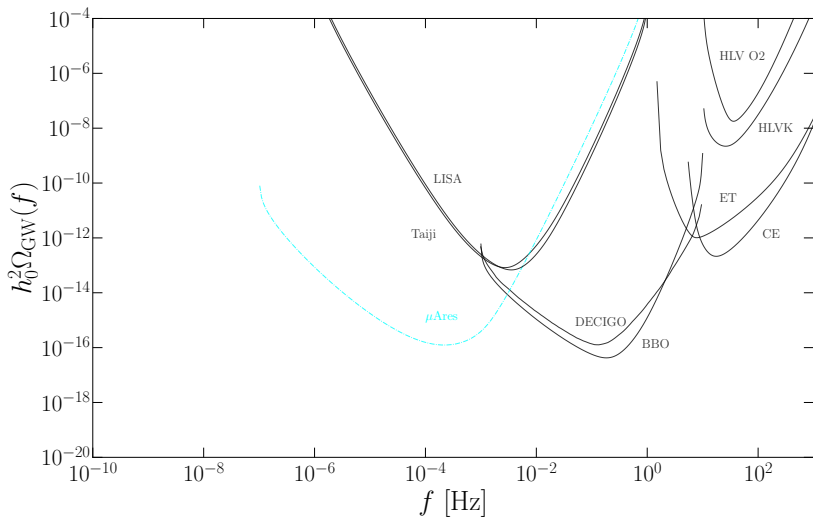


Figure: Artist's impression of LISA from Wikipedia

White paper:

[LISA Cosmology Working Group] (incl. ARP),  
*Living Rev. Rel.* **26** (2023), arXiv:2204.05434.

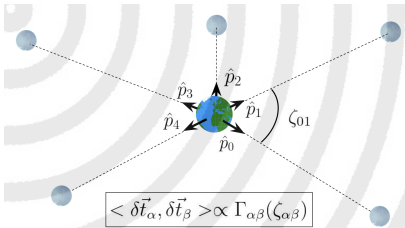
## Gravitational spectrum (space-based detectors)



## Pulsar Timing Array (PTA)

- An array of millisecond pulsars (MSP) is observed in the radio band to compute the delays on the time of arrival due to the presence of GWs.
- Collected data is the time series of residuals for each pulsar:

$$\delta t^i = t_{\text{obs}}^i - t_{\text{TM}}^i$$



Credit: Mikel Falxa

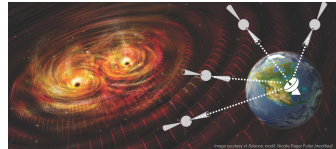
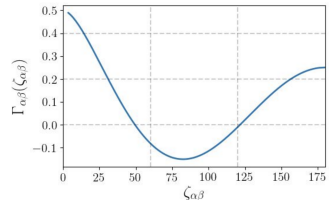


Figure: Image courtesy of Science, credit: Nicole Rager Fuller

The correlation  $\Gamma_{\alpha\beta}$  follows in GR the Hellings-Downs curve<sup>4</sup>

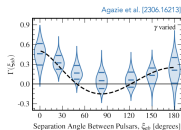


<sup>4</sup>R. W. Hellings and G. S. Downs, *Astrophys. J. Lett.* **265** (1983) L39-L42

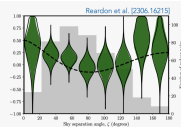
# PTA detection

- The PTA collaborations reported for the first-time evidence of a stochastic gravitational wave background on a press release on June 28, 2023.

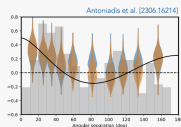
**NANOGrav:**  
68 pulsars, 16yr of data  
~3.4 $\sigma$  significance



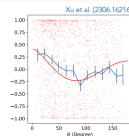
**PPTA:**  
32 pulsars, 18yr of data  
~2 $\sigma$  significance



**EPTA + InPTA:**  
25 pulsars, 24yr of data  
~3 $\sigma$  significance



**CPTA:**  
57 pulsars, 3yr of data  
~4.6 $\sigma$  significance



Credit: Andrea Mitridate

# Cosmological Gravitational Waves

- Gravity is the *weakest fundamental force*. Hence, GWs are difficult to detect but they propagate freely carrying *clean information of the source*.
- Cosmological GWs from the early Universe have the potential to provide us with *direct information on early universe physics* that is *not accessible via electromagnetic observations, complementary to collider experiments*.

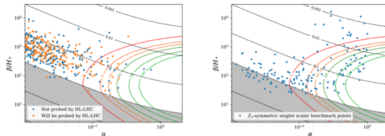
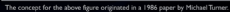


Figure 4: LEFT: Predicted values of  $\alpha$  and  $\beta/H$  for  $m_2 = 170$  GeV and 240 GeV (combined in one plot) in the general singlet model obtained by linearly varying the free parameters of the model and imposing requirements as described in the text. The mixing angles considered were  $\sin \theta = 0.01$  (blue points) and 0.1 (orange points). The models in blue (orange) are unlikely (likely) to be probed by the high-luminosity LHC. The expected LISA sensitivities correspond to  $T_* = 50$  GeV. RIGHT: Sensitivity curve for the  $Z_2$  symmetric singlet extension. In both the left and right panels we have taken  $v_w = 1$ .

From [LISA CosWG], Caprini et al., *Detecting gravitational waves from cosmological phase transitions with LISA: an update (2019)*

pre-BBN magnetic fields and gravitational wave background (GWB)



Supported by DOE

# Probing the early Universe with GWs

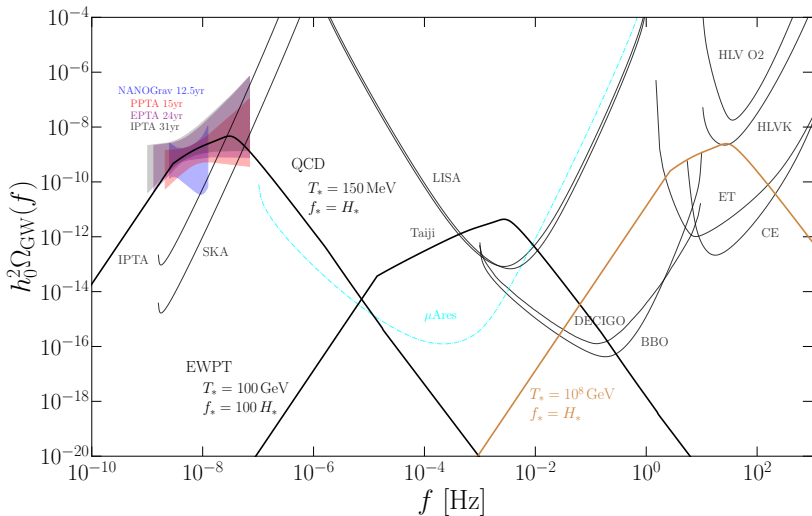
## Cosmological (pre-recombination) GW background

- Why background? Individual sources are not resolvable, superposition of single events occurring in the whole Universe.

$$f_* \simeq 1.64 \times 10^{-3} \frac{0.01}{R_* \mathcal{H}_*} \frac{T_*}{100 \text{ GeV}} \text{ Hz}$$

- Phase transitions
  - Ground-based detectors (LVK, ET, CE) frequencies are 10–1000 Hz  
Peccei-Quinn, B-L, left-right symmetries  $\sim 10^7, 10^8$  GeV.
  - Space-based detectors (**LISA**) frequencies are  $10^{-5}$ – $10^{-2}$  Hz  
**Electroweak phase transition**  $\sim 100$  GeV
  - Pulsar Timing Array (**PTA**) frequencies are  $10^{-9}$ – $10^{-7}$  Hz  
**Quark confinement (QCD) phase transition**  $\sim 100$  MeV

## Gravitational spectrum from PTs<sup>5</sup>

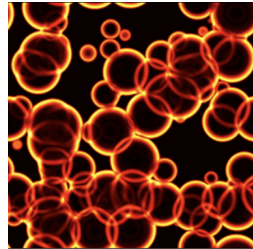
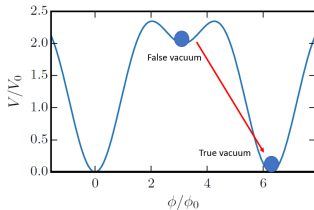


<sup>5</sup> ARP, C. Caprini, A. Neronov, D. Semikoz, *PRD* **105**, 123502 (2022)  
 A. Neronov, ARP, C. Caprini, D. Semikoz, *PRD* **103**, L041302 (2021)  
 ARP *et al.*, arXiv:2307.10744 (2023).

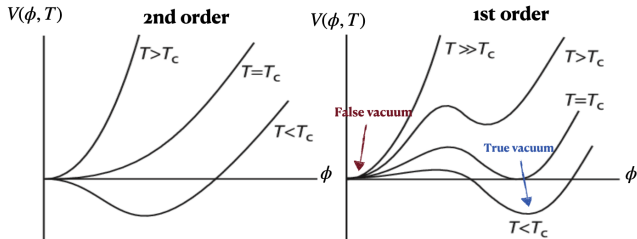


# First-order phase transition

$$V(\phi, T) = \frac{1}{2}M^2(T)\phi^2 - \frac{1}{3}\delta(T)\phi^3 + \frac{1}{4}\lambda\phi^4$$

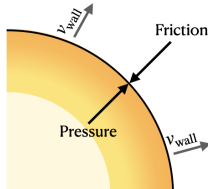


Credits: I. Stomberg



## Hydrodynamics of first-order phase transitions<sup>6</sup>

- Broken-phase bubbles are nucleated and expand
- Friction from particles yield a terminal velocity  $\xi_w$  of the bubbles
- The bubble can run away when the friction is not enough to stop the bubble's acceleration



$$\nabla_\mu T_{\text{field}}^{\mu\nu} = \frac{\partial V}{\partial \phi} \partial^\nu \phi + \eta u^\mu \partial_\mu \phi \partial^\nu \phi,$$

$$\nabla_\mu T_{\text{fluid}}^{\mu\nu} = -\frac{\partial V}{\partial \phi} \partial^\nu \phi - \eta u^\mu \partial_\mu \phi \partial^\nu \phi,$$

Boltzmann equation

$$\left( p \cdot \partial + \frac{1}{2} \partial_\mu m_X^2 \partial_\mu \right) f_X(\vec{p}, x) \delta(p^2 - m^2) = -C[f_X]$$

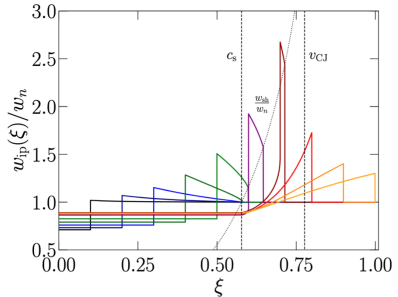
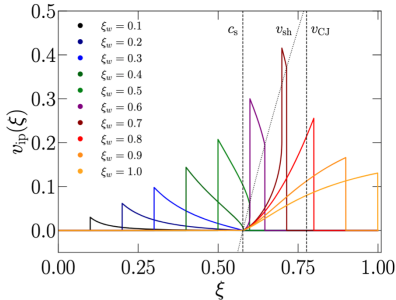
$$\partial_\mu T_{\text{fluid}}^{\mu\nu} = -\partial^\nu \phi \frac{dm^2}{d\phi} \int \frac{d^3 p}{(2\pi)^3} \frac{1}{2E_{\vec{p}}} f(p, x) \Big|_{p^0=E_{\vec{p}}}$$



$$\begin{aligned} f_{\text{eq}} &\longrightarrow \partial^\nu V_1(\phi, T) \\ \delta f &\longrightarrow -\eta \delta^\nu \phi u^\mu \partial_\mu \phi \end{aligned}$$

## Hydrodynamics of first-order phase transitions<sup>7</sup>

- Depending on the wall velocity  $\xi_w$  and the strength of the phase transition  $\alpha = \epsilon/\rho_{\text{rad}}$ , we can have 3 types of solution: *subsonic deflagrations*, *supersonic deflagrations (hybrids)* or *detonations*.

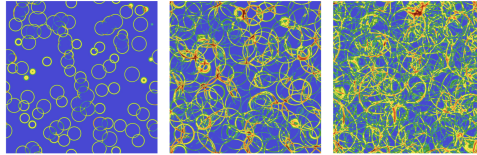


These calculations are available on CosmoGW, using  
`pip install cosmoGW`

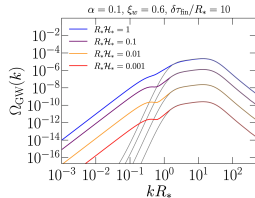
<sup>7</sup> Espinosa, Konstandin, No, Servant, *JCAP* **06** (2010) 028.

## GWs from sound waves<sup>8</sup>

- Numerical simulations of the scalar + fluid system performed by the Sussex/Helsinki group via an effective friction term indicate sound-wave regime to dominate for weak/intermediate phase transitions.



- Sound-shell model. Two scales are found that determine the GW spectrum:  $R_*$  and  $\Delta R_*$  (sound-shell thickness).



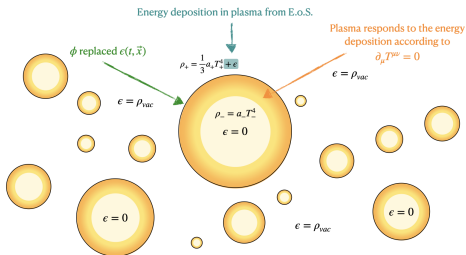
$$\Omega_{\text{GW}}(f) = 3 \tilde{\Omega}_{\text{GW}} K^2 \Upsilon(\tau_{\text{sw}})(H_* R_*) S(f R_*)$$

$$\Upsilon(\tau_{\text{sw}}) = \frac{H_* \tau_{\text{sw}}}{1 + H_* \tau_{\text{sw}}}$$

<sup>8</sup>Hindmarsh *et al.*, 2013, 2015, 2017, Cutting *et al.*, 2019, Correia *et al.*, 2025

## GWs from sound waves: Higgsless simulations<sup>9</sup>

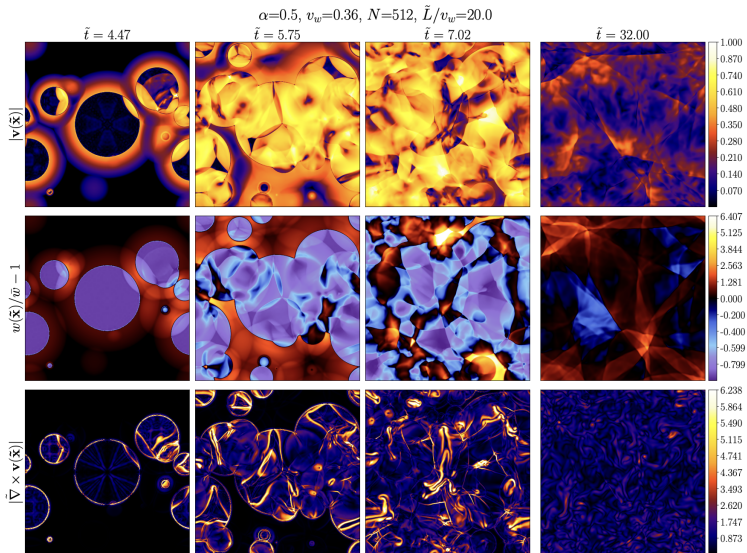
- Difficulty on simulations is due to the different scales of the scalar field  $\phi$  and the fluid shell, so one can consider a nucleation history and set the pressure and energy density by knowing the value of  $\epsilon$  and setting it during the simulation.
- Effect of bubble collisions on GWs is subdominant when sound waves are produced, so one can ignore the scalar field.
- Nucleation history is produced from an exponential probability distribution  $P(t) \propto \exp[\beta(t - t_*)]$ .



Credit: I. Stomberg

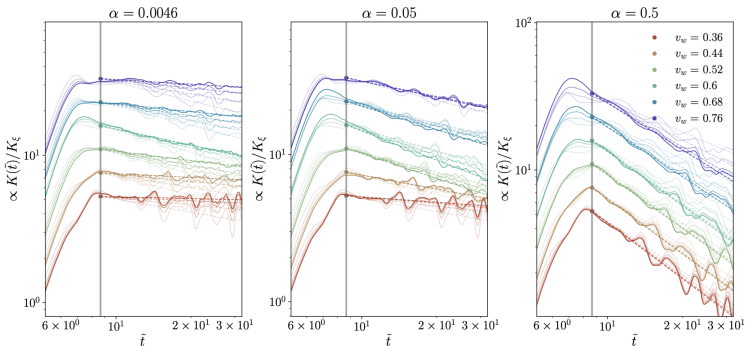
<sup>9</sup> Jinno *et al.* JCAP 02 (2023) 011, 2209.04369,  
ARP, Stomberg *et al.*, arXiv:2409.03651 (2024).

# Higgsless simulations of strong PTs<sup>10</sup>



## Higgsless simulations (results)<sup>11</sup>

- Kinetic energy decay is observed in the simulations.
- For weak and strong PTs, increasing discretization enhances the decay.
- Potential indication of the development of non-linearities (turbulence).



<sup>11</sup> ARP, Stomberg et al., arXiv:2409.03651 (2024).

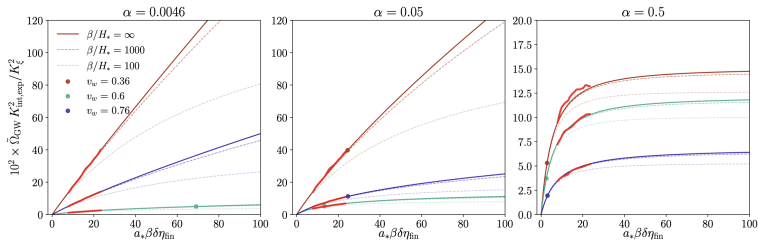
## Higgsless simulations (results)<sup>12</sup>

- In the literature, the GW spectrum from sound waves is usually assumed to be

$$\Omega_{\text{GW}}(f) = 3 \tilde{\Omega}_{\text{GW}} K^2 \mathfrak{T}(\tau_{\text{sw}}) (H_* R_*) S(f R_*)$$

- The linear growth, which only appears when expansion is neglected, is modified when the decay of the source is significant (e.g., due to the development of non-linearities).
- Extended model to proposed locally stationary UETC

$$\Omega_{\text{GW}}(f) = 3 \tilde{\Omega}_{\text{GW}} \mathfrak{K}_{\text{int,exp}}^2 (H_* R_*) S(f R_*)$$





# Higgsless simulations (results)<sup>10</sup>

- In the literature, the GW spectrum from sound waves is usually assumed to be

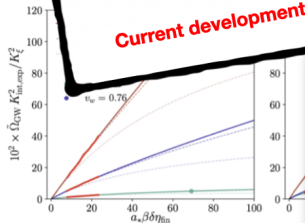
$$\Omega_{\text{GW}}(f) = 3 \tilde{\Omega}_{\text{GW}} K^2 \Upsilon(\tau_{\text{sw}}) (H_* R_*) S(\epsilon, \beta)$$

- The linear growth, which ends when the decay of the sound waves is modified

- sources based on Higgsless simulations available using public code CosmoGW

**Template for sound-shell model and template for non-linear**  
**sources based on Higgsless simulations available using public**  
**code CosmoGW**  
**(<https://github.com/cosmoGW/cosmoGW>)**  
**proceedings on arxiv with I. Stomberg (2508.04263) and tutorial**  
**available on CosmoGW documentation**

**For use: pip install cosmoGW**  
**Current developments undergoing to extend CosmoGW to**  
**other sources**



Gravitational wave spectra for cosmological phase transitions with non-linear decay of the fluid motion

Isak Stomberg<sup>a</sup> and Alberto Roper Pol<sup>b</sup>

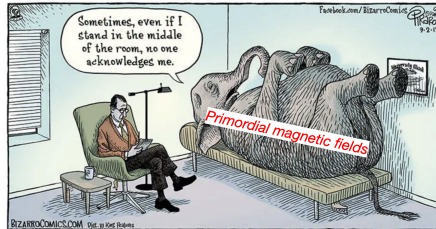
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<sup>b</sup> Département de Physique Théorique, Université de Genève, CH-1211 Genève, Switzerland



# Primordial magnetic fields

- Magnetic fields can either be produced at or present during cosmological phase transitions.
- The magnetic fields are strongly coupled to the primordial plasma and effectively produce vortical motion, inevitably leading to the development of MHD turbulence.<sup>13</sup>
- Present magnetic fields can be amplified by primordial turbulence via dynamo.<sup>14</sup>



<sup>13</sup> J. Ahonen and K. Enqvist, *Phys. Lett. B* **382**, 40 (1996).

<sup>14</sup> A. Brandenburg *et al.* (incl. ARP), *Phys. Rev. Fluids* **4**, 024608 (2019).

# Generation of primordial magnetic fields during inflation

- Magnetic fields could be amplified from quantum fluctuations during inflation.
- For inflationary magnetogenesis to be viable, conformal invariance needs to be broken,<sup>15</sup> with  $\mathcal{L} \sim f(\phi)F_{\mu\nu}F^{\mu\nu}$  or  $g(\phi)F_{\mu\nu}\tilde{F}^{\mu\nu}$ . Otherwise, magnetic field perturbations would decay with the Universe expansion.
- After reheating, magnetic fields are strongly coupled to the primordial plasma and effectively produce vortical motion, inevitably leading to the development of MHD turbulence.<sup>16</sup>
- Alternatively, small seed magnetic fields can be amplified by primordial turbulence induced by phase transitions, e.g., via dynamo.<sup>17</sup>

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<sup>15</sup> Turner & Widrow (1988), Ratra (1992), Gasperini *et al.* (1995).

<sup>16</sup> J. Ahonen and K. Enqvist, *Phys. Lett. B* **382**, 40 (1996).

<sup>17</sup> A. Brandenburg *et al.* (incl. ARP), *Phys. Rev. Fluids* **4**, 024608 (2019).

# Generation of primordial magnetic fields during phase transitions

- Parity-violating processes during the electroweak phase transition ( $T \sim 100$  GeV) are predicted by SM extensions that account for baryogenesis and can produce helical magnetic fields through sphaleron decay or B+L anomalies.<sup>18</sup>

$$\mathbf{B} = \nabla \times \mathbf{A} - i \frac{2 \sin \theta_w}{g v^2} \nabla \Phi^\dagger \times \nabla \Phi$$

- Also after inflation, axion fields can amplify and produce magnetic fields.<sup>19</sup> For example, the QCD axion could oscillate and produce magnetic fields around the QCD scale ( $T \sim 100$  MeV).<sup>20</sup>

$$\mathcal{L} \supset \frac{\phi}{f} F_{\mu\nu} \tilde{F}^{\mu\nu}$$

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<sup>18</sup> T. Vachaspati, *Phys. Rev. B* **265**, 258 (1991), T. Vachaspati, *Phys. Rev. Lett.* **87**, 251302 (2001), J. M. Cornwall, *Phys. Rev. D* **56**, 6146 (1997).

<sup>19</sup> M. M. Forbes and A. R. Zhitnitsky, *Phys. Rev. Lett.* **85**, 5268 (2000).

<sup>20</sup> Miniati, Gregori, Reville & Sarkar (2018).

# Generation of primordial magnetic fields from chiral anomalies

- Chiral magnetic effect in the early Universe at large temperatures  $T > 80$  TeV can lead to the production of hypermagnetic fields.<sup>21</sup>
- Proposed chiral magnetic effect relevant at smaller scales<sup>22</sup> leading to a chiral plasma instability and development of magnetic fields that undergo inverse cascading.<sup>23</sup>
- Chiral magnetic effect also present for zero chiral chemical potential from local inhomogeneities.<sup>24</sup>
- Chiral magnetic effect can be activated from an active source of chiral chemical potential (e.g. associated to baryogenesis at the electroweak scale)<sup>25</sup>

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<sup>21</sup> M. Joyce and M. E. Shaposhnikov, *PRL* **79**, 1193 (1997)

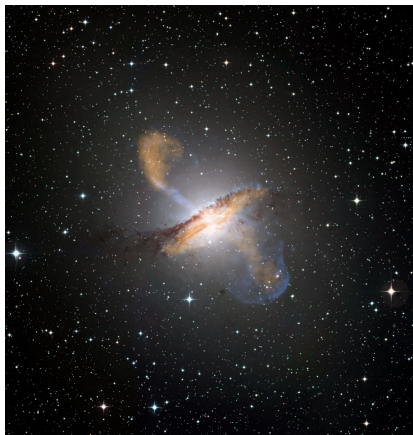
<sup>22</sup> A. Boyarsky *et al.* *PRL* **108**, 031301 (2012).

<sup>23</sup> I. Rogachevskii *et al.* *Astrophys. J. Lett* **845** (2017) L21, J. Schober *et al.* *Astrophys. J.* **858**, 124 (2018).

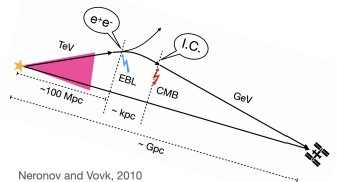
<sup>24</sup> J. Schober *et al.* *PRL* **132**, 6 (2024)

<sup>25</sup> M. Gurgunidze *et al.* (incl. ARP), arXiv:2512.09177 (2025).

## Intergalactic magnetic fields in voids of the LSS



- $\gamma$ -ray observations from distant blazars by Fermi/MAGIC collaboration show a removal of power at GeV, providing evidence for an intergalactic magnetic field.
- Lower bound of  $B \sim 10^{-16}$  G at Mpc scales.



# Relics from the early Universe?

- The observed intergalactic magnetic fields could have a **primordial** (from inflation or from a phase transition) or an **astrophysical** origin.
- Observations indicate a large volume filling factor in the voids and large correlation length scales (Mpc), favoring a primordial origin.
- Alternative possibilities to explain the blazar observations are:
  - *Beam-plasma instabilities* (suppressing the GeV emission) proposed<sup>26</sup>, but seems to not be significant at the intergalactic scales, *recently shown to be suppressed experimentally with the Super Proton Synchrotron at CERN*.<sup>27</sup>
  - *Magnetized galactic outflows*. Studied using cosmological simulations. However, it is difficult to reach large volume filling factors.<sup>28</sup>
- Primordial magnetic field present at recombination could alleviate the Hubble tension by reducing the sound horizon.<sup>29</sup>

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<sup>26</sup> Broderick et al. (2018).

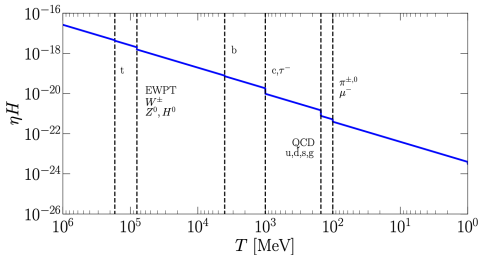
<sup>27</sup> Arrosmith et al. (2025).

<sup>28</sup> Ni et al. (2024), Tjemsland et al. (2024).

<sup>29</sup> Jedamzik & Pogosian (2020).

# GW sources in the early Universe

- Magnetohydrodynamic (MHD) sources of GWs:
  - Sound waves generated from first-order phase transitions.
  - (M)HD turbulence from first-order phase transitions.
  - Primordial magnetic fields.
- High-conductivity (low magnetic diffusivity  $\eta = 1/\sigma$ ) of the early universe leads to a high-coupling between magnetic and velocity fields.<sup>30</sup>





## Primordial magnetic fields and GWs from MHD turbulence

- During the radiation-dominated era, after they are produced, magnetic fields will decay following MHD turbulence.
- Direct numerical simulations using the `PENCIL CODE`<sup>31</sup> to solve:
  - ① Relativistic MHD equations adapted for radiation-dominated era (after electroweak symmetry is broken).
  - ② Gravitational waves equation.
- In general, large-scale simulations are necessary to solve the MHD nonlinearities (e.g., unequal-time correlators UETC and non-Gaussianities, which require simplifying assumptions in analytical studies).

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<sup>31</sup>Pencil Code Collaboration, JOSS **6**, 2807 (2020),  
<https://github.com/pencil-code/>

## MHD description

Right after the electroweak phase transition we can model the plasma using continuum MHD.

- Charge-neutral, electrically conducting fluid.
- Relativistic magnetohydrodynamic (MHD) equations.
- Radiation-dominated fluid

$$p = \rho c^2/3,$$

i.e.,  $c_s^2 = 1/3$  (ultrarelativistic EoS).

- Friedmann–Lemaître–Robertson–Walker metric

$$g_{\mu\nu} = \text{diag}\{-1, a^2, a^2, a^2\}$$

## Conservation laws for MHD turbulence

$$T^{\mu\nu}_{;\nu} = 0, \quad F^{\mu\nu}_{;\nu} = -J^\mu, \quad \tilde{F}^{\mu\nu}_{;\nu} = 0$$

In the limit of subrelativistic bulk flow:

$$\gamma^2 \sim 1 + (v/c)^2 + \mathcal{O}(v/c)^4$$

Relativistic MHD equations are reduced to<sup>32</sup>

$$\begin{aligned} \frac{\partial \ln \rho}{\partial t} &= -\frac{4}{3} (\nabla \cdot \mathbf{u} + \mathbf{u} \cdot \nabla \ln \rho) + \frac{1}{\rho} [\mathbf{u} \cdot (\mathbf{J} \times \mathbf{B}) + \eta J^2], \\ \frac{D\mathbf{u}}{Dt} &= \frac{\mathbf{u}}{3} (\nabla \cdot \mathbf{u} + \mathbf{u} \cdot \nabla \ln \rho) - \frac{\mathbf{u}}{\rho} [\mathbf{u} \cdot (\mathbf{J} \times \mathbf{B}) + \eta J^2] \\ &\quad - \frac{1}{4} \nabla \ln \rho + \frac{3}{4\rho} \mathbf{J} \times \mathbf{B} + \frac{2}{\rho} \nabla \cdot (\rho \nu \mathbf{S}), \\ \frac{\partial \mathbf{B}}{\partial t} &= \nabla \times (\mathbf{u} \times \mathbf{B} - \eta \mathbf{J}), \quad \mathbf{J} = \nabla \times \mathbf{B}, \end{aligned} \tag{1}$$

for a flat expanding universe with comoving and normalized

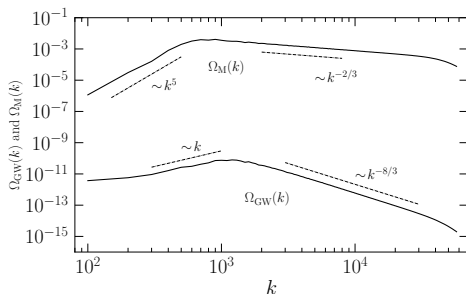
$\rho = a^4 \rho_{\text{phys}}$ ,  $\rho = a^4 \rho_{\text{phys}}$ ,  $B_i = a^2 B_{i,\text{phys}}$ ,  $u_i$ , and conformal time  $t$  ( $dt = a dt_c$ ).

<sup>32</sup>A. Brandenburg, *et al.*, *Phys. Rev. D* **54**, 1291 (1996).

ARP, Midiri, *Relativistic Magnetohydrodynamics in the early Universe*, accepted in JCAP, arXiv:2501.05732 (2025).

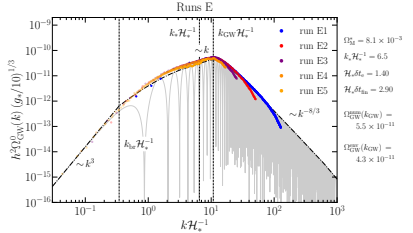
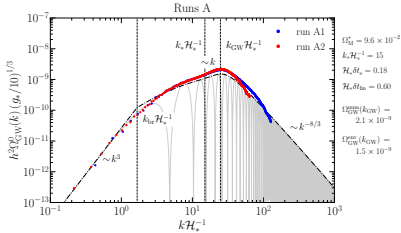
## Numerical results for decaying MHD turbulence<sup>33</sup>

$$1152^3, k_* = 2\pi \times 100, \Omega_M \sim 10^{-2}, \sigma_M = 1$$



- **Characteristic  $k$  scaling in the subinertial range for the GW spectrum.**
- $k^2$  expected at scales  $k < k_*$  and  $k^3$  at  $k < H_*$  according to the “top-hat” model (Caprini *et al.*, 2020).

# Numerical results for decaying MHD turbulence<sup>34</sup>



| run | $\Omega_{\text{M}}^*$ | $k_* \mathcal{H}_*^{-1}$ | $\mathcal{H}_* \delta t_e$ | $\mathcal{H}_* \delta t_{\text{fin}}$ | $\Omega_{\text{GW}}^{\text{num}}(k_{\text{GW}})$ | $[\Omega_{\text{GW}}^{\text{env}}/\Omega_{\text{GW}}^{\text{num}}](k_{\text{GW}})$ | $n$ | $\mathcal{H}_* L$ | $\mathcal{H}_* t_{\text{end}}$ | $\mathcal{H}_* \eta$ |
|-----|-----------------------|--------------------------|----------------------------|---------------------------------------|--------------------------------------------------|------------------------------------------------------------------------------------|-----|-------------------|--------------------------------|----------------------|
| A1  | $9.6 \times 10^{-2}$  | 15                       | 0.176                      | 0.60                                  | $2.1 \times 10^{-9}$                             | 1.357                                                                              | 768 | $6\pi$            | 9                              | $10^{-7}$            |
| A2  | —                     | —                        | —                          | —                                     | —                                                | —                                                                                  | 768 | $12\pi$           | 9                              | $10^{-6}$            |
| E1  | $8.1 \times 10^{-3}$  | 6.5                      | 1.398                      | 2.90                                  | $5.5 \times 10^{-11}$                            | 1.184                                                                              | 512 | $4\pi$            | 8                              | $10^{-7}$            |
| E2  | —                     | —                        | —                          | —                                     | —                                                | —                                                                                  | 512 | $10\pi$           | 18                             | $10^{-7}$            |
| E3  | —                     | —                        | —                          | —                                     | —                                                | —                                                                                  | 512 | $20\pi$           | 61                             | $10^{-7}$            |
| E4  | —                     | —                        | —                          | —                                     | —                                                | —                                                                                  | 512 | $30\pi$           | 114                            | $10^{-7}$            |
| E5  | —                     | —                        | —                          | —                                     | —                                                | —                                                                                  | 512 | $60\pi$           | 234                            | $10^{-7}$            |

## Analytical model for GWs from decaying turbulence

- Assumption: magnetic or velocity field evolution  $\delta t_e \sim 1/(u_* k_*)$  is slow compared to the GW dynamics ( $\delta t_{\text{GW}} \sim 1/k$ ) at all  $k \gtrsim u_* k_*$ .
- We can derive an analytical expression for nonhelical fields of the envelope of the oscillations<sup>35</sup> of  $\Omega_{\text{GW}}(k)$ .

$$\Omega_{\text{GW}}(k, t_{\text{fin}}) \approx 3 \left( \frac{k}{k_*} \right)^3 \Omega_{\text{M}}^{*2} \frac{\mathcal{C}(\alpha)}{\mathcal{A}^2(\alpha)} p_{\Pi} \left( \frac{k}{k_*} \right) \\ \times \begin{cases} \ln^2[1 + \mathcal{H}_* \delta t_{\text{fin}}] & \text{if } k \delta t_{\text{fin}} < 1, \\ \ln^2[1 + (k/\mathcal{H}_*)^{-1}] & \text{if } k \delta t_{\text{fin}} \geq 1. \end{cases}$$

- $p_{\Pi}$  is the anisotropic stress spectrum and depends on spectral shape, can be approximated for a von Kármán spectrum as<sup>36</sup>

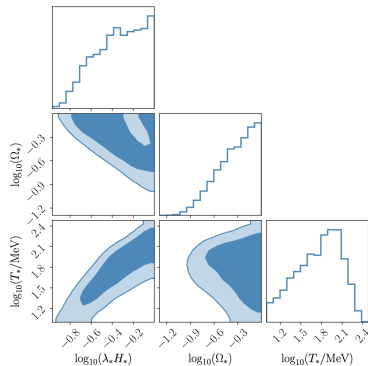
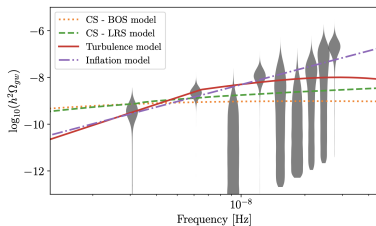
$$p_{\Pi}(k/k_*) \simeq \left[ 1 + \left( \frac{k}{2.2k_*} \right)^{2.15} \right]^{-11/(3 \times 2.15)}$$

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<sup>35</sup> ARP *et al.*, *Phys. Rev. D* **105**, 123502 (2022).

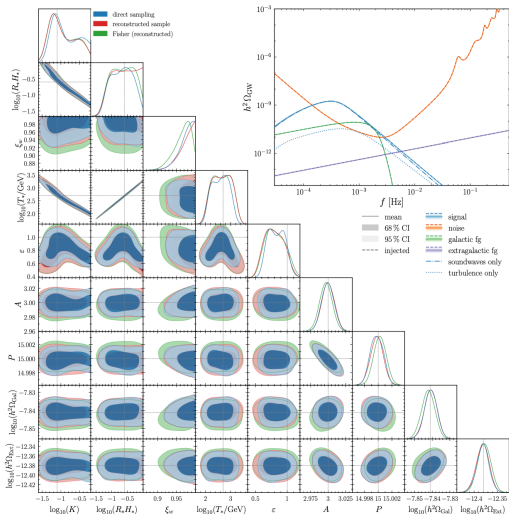
<sup>36</sup> ARP *et al.*, arXiv:2307.10744 (2023).

## Primordial magnetic fields constraints with PTA<sup>37</sup>



<sup>37</sup>[EPTA collab.] (incl. ARP), arXiv:2306.16227 (2023).

# Parameter reconstruction of GW signal for sound waves and turbulence from a PT with LISA<sup>38</sup>





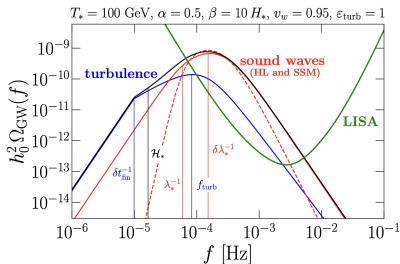
# GW sources in the early Universe

- Magnetohydrodynamic (MHD) sources of GWs:
  - Sound waves generated from first-order phase transitions.
  - (M)HD turbulence from first-order phase transitions.
  - Primordial magnetic fields.

- Other sources of GWs include

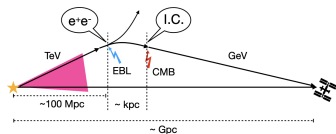
- Bubble collisions.
- Cosmic strings.
- Primordial black holes.
- Inflation.

ARP *et al.*, 2307.10744, 2308.12943

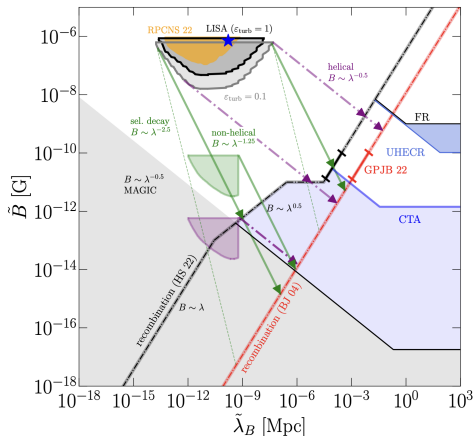


# Multi-messenger studies of primordial magnetic fields<sup>30</sup>

- Primordial magnetic fields would evolve through the history of the universe up to the present time and could explain the lower bounds in cosmic voids derived by the Fermi collaboration.<sup>31</sup>



- Maximum amplitude of primordial magnetic fields is constrained by the big bang nucleosynthesis.<sup>32</sup>
- Additional constraints from CMB, Faraday Rotation, ultra-high energy cosmic rays (UHECR).



<sup>30</sup> ARP et al., arXiv:2307.10744 (2023).

<sup>31</sup> A. Neronov and I. Vovk, *Science* **328**, 73 (2010).

<sup>32</sup> V. F. Shvartsman, *Pisma Zh. Eksp. Teor. Fiz.* **9**, 315 (1969).

# Conclusions

- Velocity and magnetic fields in the early universe can significantly contribute to the cosmological GW background (GWB) via sound waves and (M)HD turbulence.
- An observation of a GWB from the early Universe could probe BSM physics that are not accessible by other means or that are complementary to collider experiments.
- The GWB produced by non-linear fluid (and magnetic field) dynamics requires, in general, performing high-resolution numerical simulations, which can be done using open-source codes like `PENCIL CODE` or `COSMOLATTICE`.
- We expect the dynamics of first-order phase transitions to be affected by gauge fields, with its impact on the resulting GWB and on the fluid dynamics sourcing GWs being at the moment uncertain.
- Magnetic fields are ubiquitous in nature at all scales, from the smallest structures up to galaxies, cluster of galaxies, filaments, and potentially in cosmic voids of the LSS.
- Primordial magnetic fields produced in the early Universe can be probed by other cosmological (CMB) and astrophysical ( $\gamma$ -ray, radio observations at voids, filaments of the LSS), providing us a multi-messenger approach to study early Universe physics.
- To extract clean information from a GWB observation in LISA, PTA, or ET, we will need to have accurate modeling of the resulting GWB, since the reconstruction of the signal is very challenging (multiple cosmo sources, astrophysical background, detectors noise).



# Thank You!

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