Multi-messenger searches of primordial magnetic fields and gravitational waves

Max Planck Institute for Radioastronomy (MPifR) Colloquium
(Dec. 5, 2025. Bonn, Germany)

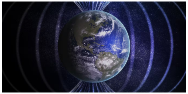
Max-Planck-Institut
für Radioastronomie

Alberto Roper Pol
University of Geneva

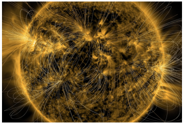
SNSF Ambizione grant (2023–2027): "*Exploring the early universe with gravitational waves and primordial magnetic fields.*"

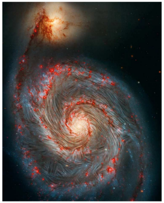
Collaborators: A. Brandenburg (Nordita), C. Caprini (CERN), T. Kahniashvili (CMU), A. Kosowsky (PittU), S. Mandal (SBU), A. Neronov (APC), D. Semikoz (APC)

Our magnetized Universe

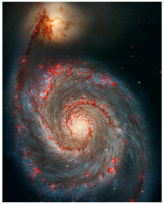


$\sim 1\,\mathrm{G}$





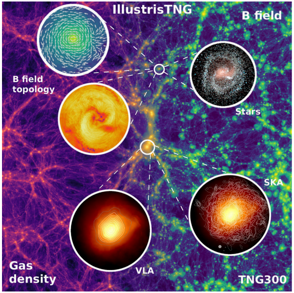
(FIR 145 μm)



(Radio 6 cm Fletcher et. al)

$\sim 10^{-5}\,\mathrm{G}$

Borlaff et al. (2021), SOFIA, M51

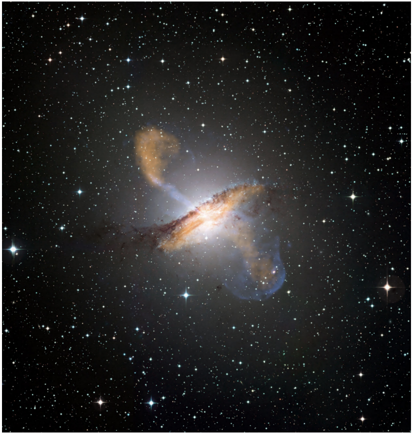


Marinacci et al. (2017), Illustris TNG

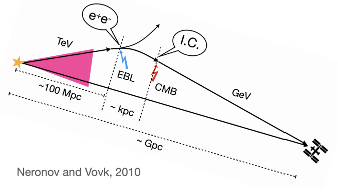
$\sim 10^{-9} - 10^{-6}\,\mathrm{G}$

- Dense regions in our Universe are strongly magnetized with $B \sim 10\mu$G in galaxies and clusters ($\sim$10 kpc). Radio and far-infrared (FIR) observations.
- Filaments also contain magnetic fields with $B \sim$10 nG (LOFAR observations of Faraday rotation measurements).
- **Past colloquium on Nov. 21: "The magnetic history of our Universe" by J. Schober**

Intergalactic magnetic fields in voids of the LSS



- γ-ray observations from distant blazars by Fermi/MAGIC collaboration show a removal of power at GeV, providing evidence for an intergalactic magnetic field.
- Lower bound of $B \sim 10^{-16}$ G at Mpc scales.
- **Past colloquium on Sep. 26: "Intergalactic and Cosmological Magnetic Fields" by A. Neronov**



Neronov and Vovk, 2010

Relics from the early Universe?

- The observed intergalactic magnetic fields could have a **primordial** (from inflation or from a phase transition) or an **astrophysical** origin.
- Observations indicate a large volume filling factor in the voids and large correlation length scales (Mpc), favoring a primordial origin.
- Alternative possibilities to explain the blazar observations are:
    - *Beam-plasma instabilities* (suppressing the GeV emission) proposed[1], but seems to not be significant at the intergalactic scales, *recently shown to be suppressed experimentally with the Super Proton Synchrotron at CERN.*[2]
    - *Magnetized galactic outflows.* Studied using cosmological simulations. However, it is difficult to reach large volume filling factors.[3]
- Primordial magnetic field present at recombination could alleviate the Hubble tension by reducing the sound horizon.[4]

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[1] Broderick et al. (2018).

[2] Arroswmith et al. (2025).

[3] Ni et al. (2024), Tjemsland et al. (2024).

[4] Jedamzik & Pogosian (2020).

Generation of primordial magnetic fields during inflation

- Magnetic fields could be amplified from quantum fluctuations during inflation or during cosmological phase transitions.
- For inflationary magnetogenesis to be viable, conformal invariance needs to be broken,[5] with $\mathcal{L} \sim f(\phi) F_{\mu\nu} F^{\mu\nu}$ or $g(\phi) F_{\mu\nu} \tilde{F}^{\mu\nu}$. Otherwise, magnetic field perturbations would decay with expansion.
- After reheating, magnetic fields are strongly coupled to the primordial plasma and effectively produce vortical motion, inevitably leading to the development of MHD turbulence.[6]
- Alternatively, small seed magnetic fields can be amplified by primordial turbulence induced by phase transitions, e.g., via dynamo.[7]
- Inhomogeneities in the Higgs field in low-scale electroweak hybrid inflation.[8]

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[5] Turner & Widrow (1988), Ratra (1992), Gasperini *et al.* (1995).

[6] J. Ahonen and K. Enqvist, *Phys. Lett. B* **382**, 40 (1996).

[7] A. Brandenburg *et al.* (incl. ARP), *Phys. Rev. Fluids* **4**, 024608 (2019).

[8] M. Joyce and M. E. Shaposhnikov, *Phys. Rev. Lett.* **79**, 1193 (1997), J. García-Bellido *et al.*, *Phys. Rev. D* **60**, 123504 (1999).

Generation of primordial magnetic fields during phase transitions

- Parity-violating processes during the electroweak phase transition ($T \sim 100$ GeV) are predicted by SM extensions that account for baryogenesis and can produce helical magnetic fields through sphaleron decay or B+L anomalies.[9]

$$\boldsymbol{B} = \boldsymbol{\nabla} \times \boldsymbol{A} - i\frac{2\sin\theta_w}{gv^2}\boldsymbol{\nabla}\Phi^\dagger \times \boldsymbol{\nabla}\Phi$$

- Also after inflation, axion fields can amplify and produce magnetic fields.[10] For example, the QCD axion could oscillate and produce magnetic fields around the QCD scale ($T \sim 100$ MeV).[11]

$$\mathcal{L} \supset \frac{\phi}{f} F_{\mu\nu}\tilde{F}^{\mu\nu}$$

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9. T. Vachaspati, *Phys. Rev. B* **265**, 258 (1991), T. Vachaspati, *Phys. Rev. Lett.* **87**, 251302 (2001), J. M. Cornwall, *Phys. Rev. D* **56**, 6146 (1997).

10. M. M. Forbes and A. R. Zhitnitsky, *Phys. Rev. Lett.* **85**, 5268 (2000).

11. Miniati, Gregori, Reville & Sarkar (2018).

Generation of primordial magnetic fields from chiral anomalies

- Chiral magnetic effect in the early Universe at large temperatures $T > 80$ TeV can lead to the production of hypermagnetic fields.[12]
- Proposed chiral magnetic effect relevant at smaller scales[13] leading to a chiral plasma instability and development of magnetic fields that undergo inverse cascading.[14]
- Chiral magnetic effect also present for zero chiral chemical potential from local inhomogeneities.[15]
- Chiral magnetic effect can be activated from an active source of chiral chemical potential (e.g. associated to baryogenesis as the electroweak scale)[16]

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[12] M. Joyce and M. E. Shaposhnikov, *PRL* **79**, 1193 (1997)

[13] A. Boyarsky *et al.* *PRL* **108**, 031301 (2012).

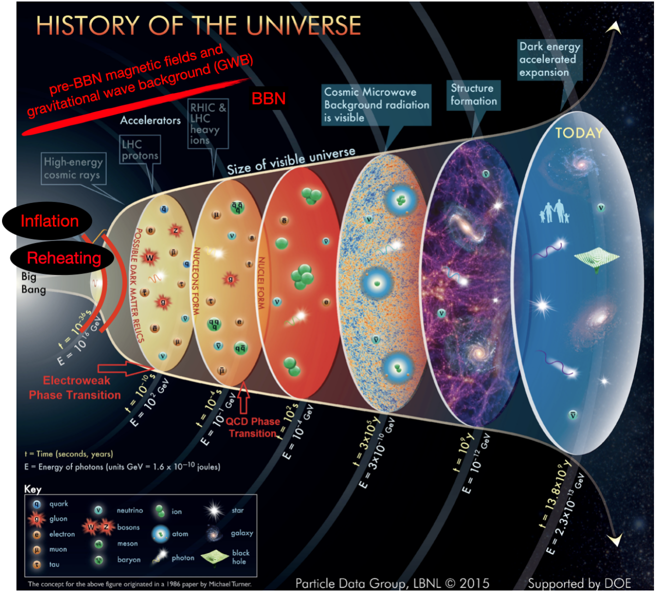
[14] I. Rogachevskii *et al.* *Astrophys. J. Lett* **845** (2017) L21, J. Schober *et al.* *Astrophy. J.* **858**, 124 (2018).

[15] J. Schober *et al.* *PRL* **132**, 6 (2024)

[16] M. Gurgenidze *et al.* (incl. ARP), *in preparation* (2025).

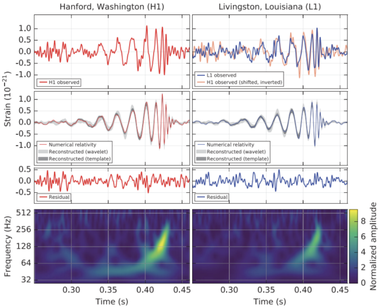
GWs from the early Universe

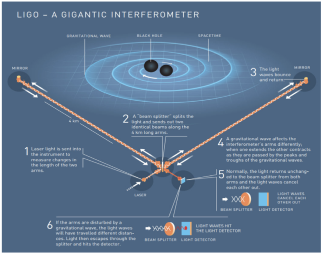
- Gravity is the *weakest fundamental force*. Hence, GWs are difficult to detect but they propagate freely carrying *clean information of the source.*
- Primordial magnetic fields and MHD turbulence in the primordial plasma can produce *a gravitational wave background* potentially observable with *LISA, PTA, or other GW experiments*.
- Observations of gravitational wave backgrounds and primordial magnetic fields can be combined for *multi-messenger studies of the primordial Universe.*
- GWs from the early Universe have the potential to provide us with *direct information on early universe physics* that is *not accessible via electromagnetic observations, complementary to collider experiments*.


HISTORY OF THE UNIVERSE
pre-BBN magnetic fields and gravitational wave background (GWB)
BBN
Accelerators
RHIC & LHC heavy ions
LHC protons
High-energy cosmic rays
Cosmic Microwave Background radiation is visible
Structure formation
Dark energy accelerated expansion
TODAY
Size of visible universe
Inflation
Reheating
Big Bang
POSSIBLE DARK MATTER RELICS
NUCLEONS FORM
NUCLEI FORM
t = 10^-36 s
E = 10^16 GeV
Electroweak Phase Transition
t = 10^-10 s
E = 10^2 GeV
t = 10^-4 s
E = 10^-1 GeV
QCD Phase Transition
t = 10^2 s
E = 10^-4 GeV
t = 3x10^5 y
E = 3x10^-10 GeV
t = 10^9 y
E = 10^-12 GeV
t = 13.8x10^9 y
E = 2.3x10^-13 GeV
t = Time (seconds, years)
E = Energy of photons (units GeV = 1.6 x 10^-10 joules)
Key
quark
gluon
electron
muon
tau
neutrino
bosons
meson
baryon
ion
atom
photon
star
galaxy
black hole
The concept for the above figure originated in a 1986 paper by Michael Turner.
Particle Data Group, LBNL © 2015
Supported by DOE

First detections of gravitational waves

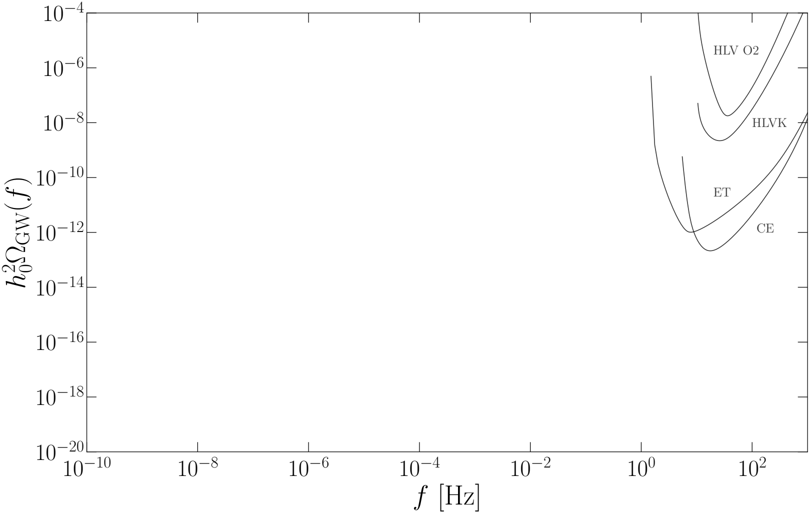
- Gravitational waves are opening a new window into our understanding of the Universe
  - First event GW150914 detected by LIGO-Virgo collaboration[17]





[17] [LIGO-Virgo Collaboration], *Phys. Rev. Lett.* **116**, 061102 (2016)

Gravitational spectrum (ground-based detectors)



LISA

- Laser Interferometer Space Antenna (LISA) is a space–based GW detector
- Approved in 2017 as one of the main research missions of ESA (L3) with NASA collaboration
- Launch planned for 2035
- Composed by three spacecrafts in a distance of 2.5M km
- LISA cosmology working group (since 2015, $\sim$ 230 members)

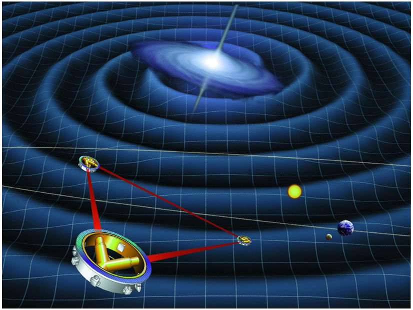


Figure: Artist's impression of LISA from Wikipedia

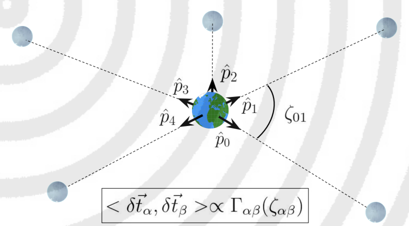
White paper:

[LISA Cosmology Working Group] (incl. ARP), *Living Rev. Rel.* **26** (2023), arXiv:2204.05434.

Pulsar Timing Array (PTA)

- An array of millisecond pulsars (MSP) is observed in the radio band to compute the delays on the time of arrival due to the presence of GWs.
- Collected data is the time series of residuals for each pulsar:

$$\delta t^i = t^i_{\mathrm{obs}} - t^i_{\mathrm{TM}}$$



Credit: Mikel Falxa

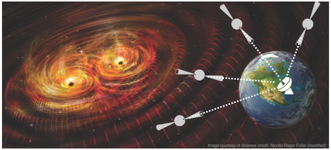
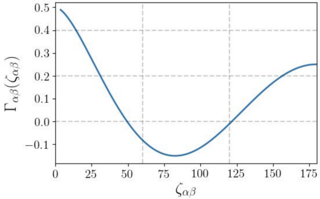


Figure: Image courtesy of Science, credit: Nicolle Rager Fuller

The correlation $\Gamma_{\alpha\beta}$ follows in GR the Hellings-Downs curve[18]



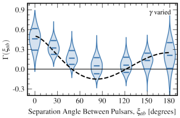
[18] R. W. Hellings and G. S. Downs, *Astrophys. J. Lett.* **265** (1983) L39-L42.

PTA detection

- The PTA collaborations reported for the first-time evidence of a stochastic gravitational wave background on a press release on June 28, 2023 (plus a series of papers by each collaboration).

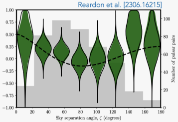
NANOGrav:
68 pulsars, 16yr of data
~3-4σ significance

Agazie et al. [2306.16213]



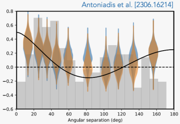
PPTA:
32 pulsars, 18yr of data
~2σ significance

Reardon et al. [2306.16215]



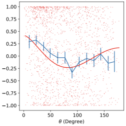
EPTA + InPTA:
25 pulsars, 24yr of data
~3σ significance

Antoniadis et al. [2306.16214]



CPTA:
57 pulsars, 3yr of data
~4.6σ significance

Xu et al. [2306.16216]



Credit: Andrea Mitridate

Probing the early Universe with GWs

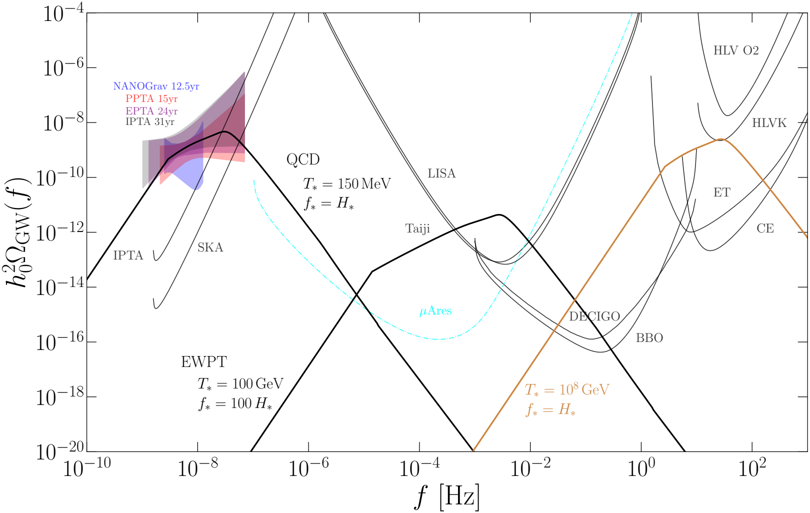
Cosmological (pre-recombination) GW background

- Why background? Individual sources are not resoluble, superposition of single events occurring in the whole Universe.

$$f_* \simeq 1.64 \times 10^{-3} \frac{100}{R_* \mathcal{H}_*} \frac{T_*}{100\,\mathrm{GeV}}\,\mathrm{Hz}$$

- Phase transitions
  - Ground-based detectors (LVK, ET, CE) frequencies are 10–1000 Hz
    Peccei-Quinn, B-L, left-right symmetries $\sim 10^7, 10^8$ GeV.
  - Space-based detectors **(LISA)** frequencies are 10^{-5}–10^{-2} Hz
    Electroweak phase transition $\sim$ 100 GeV
  - Pulsar Timing Array **(PTA)** frequencies are 10^{-9}–10^{-7} Hz
    Quark confinement (QCD) phase transition $\sim$ 100 MeV

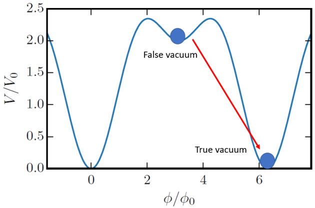
Gravitational spectrum (turbulence from PTs)[19]

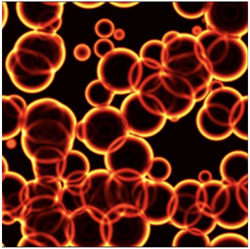


[19] ARP, C. Caprini, A. Neronov, D. Semikoz, *PRD* **105**, 123502 (2022)
A. Neronov, ARP, C. Caprini, D. Semikoz, *PRD* **103**, L041302 (2021)
ARP *et al.*, arXiv:2307.10744 (2023).

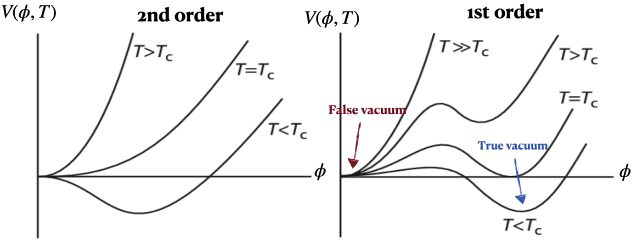
First–order phase transition

$$V(\phi, T) = \frac{1}{2}M^2(T)\phi^2 - \frac{1}{3}\delta(T)\phi^3 + \frac{1}{4}\lambda\phi^4$$



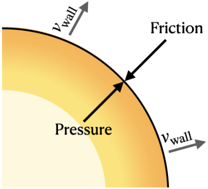


Credits: I. Stomberg



Hydrodynamics of first-order phase transitions[20]

- Broken-phase bubbles are nucleated and expand
- Friction from particles yield a terminal velocity ξ_w of the bubbles
- The bubble can run away when the friction is not enough to stop the bubble's acceleration

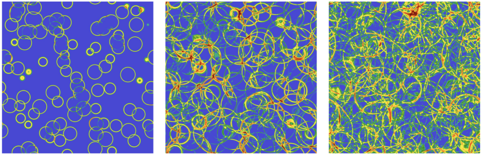


$$\nabla_\mu T^{\mu\nu}_{\text{field}} = \frac{\partial V}{\partial \phi}\partial^\nu \phi + \eta u^\mu \partial_\mu \phi \partial^\nu \phi \,,$$
$$\nabla_\mu T^{\mu\nu}_{\text{fluid}} = -\frac{\partial V}{\partial \phi}\partial^\nu \phi - \eta u^\mu \partial_\mu \phi \partial^\nu \phi \,,$$

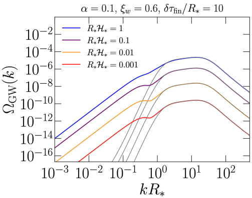
[20] Espinosa, Konstandin, No, Servant, *JCAP* **06** (2010) 028.

GWs from sound waves[21]

- Numerical simulations of the scalar + fluid system performed by the Sussex/Helsinki group via an effective friction term indicate sound-wave regime to dominate for weak/intermediate phase transitions.



- Sound-shell model. Two scales are found that determine the GW spectrum: R_* and ΔR_* (sound-shell thickness).



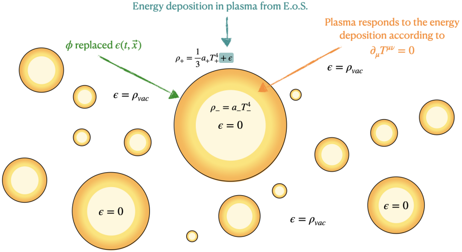
$$\Omega_{\rm GW}(f) = 3\,\tilde{\Omega}_{\rm GW}\,K^2\,\Upsilon(\tau_{\rm sw})(H_* R_*)\,S(f\,R_*)$$

$$\Upsilon(\tau_{\rm sw}) = \frac{H_*\tau_{\rm sw}}{1 + H_*\tau_{\rm sw}}$$

[21] Hindmarsh *et al.*, 2013, 2015, 2017, Cutting *et al.*, 2019, Correia *et al.*, 2025
Hindmarsh, 2016, Hindmarsh, Hijazi, 2019, ARP, Procacci, Caprini, 2023.

GWs from sound waves: Higgsless simulations[22]

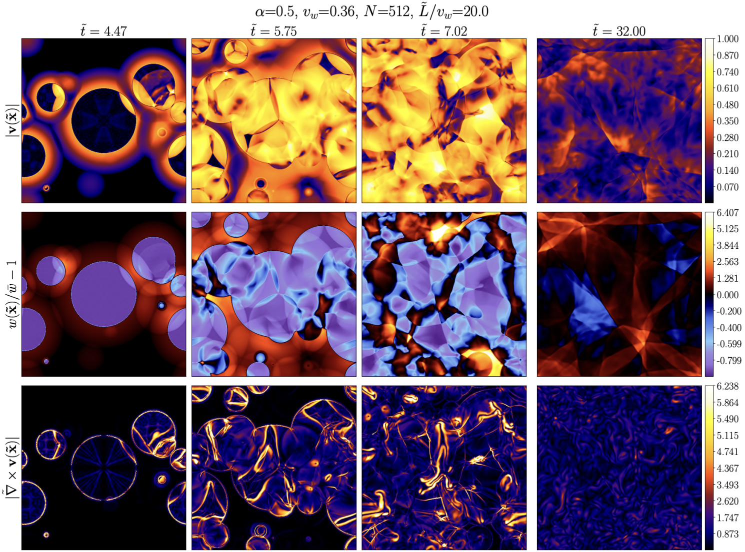
- Difficulty on simulations is due to the different scales of the scalar field ϕ and the fluid shell, so one can consider a nucleation history and set the pressure and energy density by knowing the value of ϵ and setting it during the simulation.
- Effect of bubble collisions on GWs is subdominant when sound waves are produced, so one can ignore the scalar field.
- Nucleation history is produced from an exponential probability distribution $P(t) \propto \exp\left[\beta(t - t_*)\right]$.



Credit: I. Stomberg

[22] Jinno *et al.* JCAP 02 (2023) 011, 2209.04369,
ARP, Stomberg et al., arXiv:2409.03651 (2024).

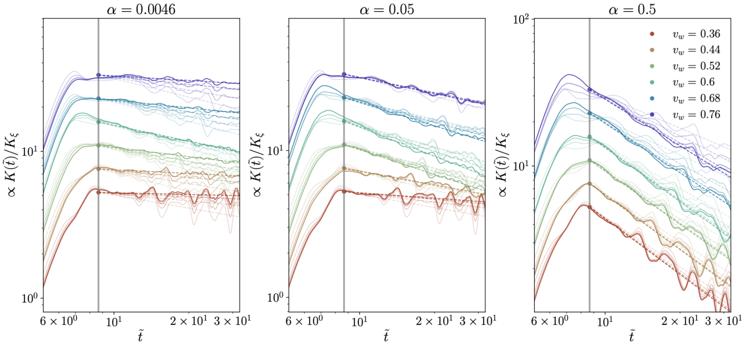
Higgsless simulations of strong PTs[23]



[23] ARP, Stomberg et al., arXiv:2409.03651.

Higgsless simulations (results)[24]

- Kinetic energy decay is observed in the simulations.
- For weak and strong PTs, increasing discretization enhances the decay.
- Potential indication of the development of non-linearities (turbulence).



[24] ARP, Stomberg et al., arXiv:2409.03651 (2024).

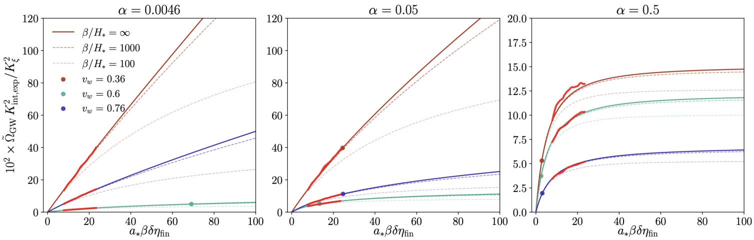
Higgsless simulations (results)[25]

- In the literature, the GW spectrum from sound waves is usually assumed to be

$$\Omega_{\rm GW}(f) = 3\,\tilde{\Omega}_{\rm GW}\, K^2\, \Upsilon(\tau_{\rm sw})\,(H_* R_*)\, S(f\, R_*)$$

- The linear growth, which only appears when expansion is neglected, is modified when the decay of the source is significant (e.g., due to the development of non-linearities).
- Extended model to proposed locally stationary UETC

$$\Omega_{\rm GW}(f) = 3\,\tilde{\Omega}_{\rm GW}\, \mathbf{K}^2_{\rm int,exp}\,(H_* R_*)\, S(f\, R_*)$$



[25] ARP, Stomberg et al., arXiv:2409.03651 (2024).

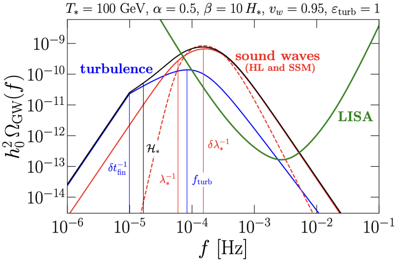
GW sources in the early Universe

- Magnetohydrodynamic (MHD) sources of GWs:
  - Sound waves generated from first-order phase transitions.
  - (M)HD turbulence from first-order phase transitions.
  - Primordial magnetic fields.

- High-conductivity of the early universe leads to a high-coupling between magnetic and velocity fields.

- Other sources of GWs include
  - Bubble collisions.
  - Cosmic strings.
  - Primordial black holes.
  - Inflation.

ARP *et al.*, 2307.10744, 2308.12943



Primordial magnetic fields and GWs from MHD turbulence

- During the radiation-dominated era, after they are produced, magnetic fields will decay following MHD turbulence.
- Direct numerical simulations using the PENCIL CODE[26] to solve:
  1. Relativistic MHD equations adapted for radiation-dominated era (after electroweak symmetry is broken).
  2. Gravitational waves equation.
- In general, large-scale simulations are necessary to solve the MHD nonlinearities (e.g., unequal-time correlators UETC and non-Gaussianities, which require simplifying assumptions in analytical studies).

[26] Pencil Code Collaboration, JOSS **6**, 2807 (2020), `https://github.com/pencil-code/`

MHD description

Right after the electroweak phase transition we can model the plasma using continuum MHD.

- Charge-neutral, electrically conducting fluid.
- Relativistic magnetohydrodynamic (MHD) equations.
- Radiation-dominated fluid

$$p = \rho c^2/3,$$

  i.e., $c_s^2 = 1/3$ (ultrarelativistic EoS).
- Friedmann–Lemaître–Robertson–Walker metric

$$g_{\mu\nu} = \mathrm{diag}\{-1, a^2, a^2, a^2\}$$

Conservation laws for MHD turbulence

$$T^{\mu\nu}{}_{;\nu} = 0, \quad F^{\mu\nu}{}_{;\nu} = -J^{\mu}, \quad \tilde{F}^{\mu\nu}{}_{;\nu} = 0$$

In the limit of subrelativistic bulk flow:

$$\gamma^2 \sim 1 + (v/c)^2 + \mathcal{O}(v/c)^4$$

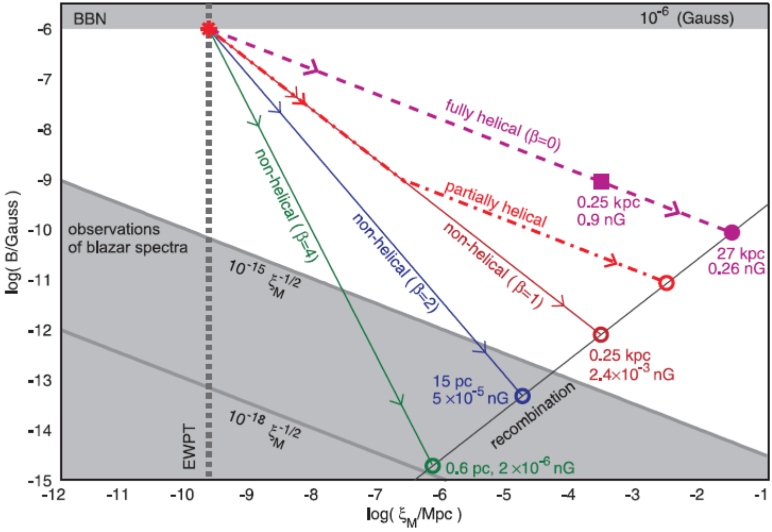
Relativistic MHD equations are reduced to[27]

$$\begin{aligned}
\frac{\partial \ln \rho}{\partial t} &= -\frac{4}{3}\left(\boldsymbol{\nabla}\cdot\boldsymbol{u} + \boldsymbol{u}\cdot\boldsymbol{\nabla}\ln\rho\right) + \frac{1}{\rho}\left[\boldsymbol{u}\cdot(\boldsymbol{J}\times\boldsymbol{B}) + \eta\boldsymbol{J}^2\right], \\
\frac{D\boldsymbol{u}}{Dt} &= \frac{\mathbf{u}}{3}\left(\boldsymbol{\nabla}\cdot\boldsymbol{u} + \boldsymbol{u}\cdot\boldsymbol{\nabla}\ln\rho\right) - \frac{\boldsymbol{u}}{\rho}\left[\boldsymbol{u}\cdot(\boldsymbol{J}\times\boldsymbol{B}) + \eta J^2\right] \\
&\quad - \frac{1}{4}\boldsymbol{\nabla}\ln\rho + \frac{3}{4\rho}\boldsymbol{J}\times\boldsymbol{B} + \frac{2}{\rho}\boldsymbol{\nabla}\cdot(\rho\nu\boldsymbol{S}), \\
\frac{\partial\boldsymbol{B}}{\partial t} &= \boldsymbol{\nabla}\times(\boldsymbol{u}\times\boldsymbol{B} - \eta\boldsymbol{J}), \quad \boldsymbol{J} = \boldsymbol{\nabla}\times\boldsymbol{B},
\end{aligned} \tag{1}$$

for a flat expanding universe with comoving and normalized $p = a^4 p_{\rm phys}, \rho = a^4 \rho_{\rm phys}, B_i = a^2 B_{i,_{\rm phys}}, u_i$, and conformal time t $({\rm d}t = a{\rm d}t_c)$.

[27] A. Brandenburg, *et al.*, *Phys. Rev. D* **54**, 1291 (1996).
ARP, Midiri, *Relativistic Magnetohydrodynamics in the early Universe*, arXiv:2501.05732 (2025).

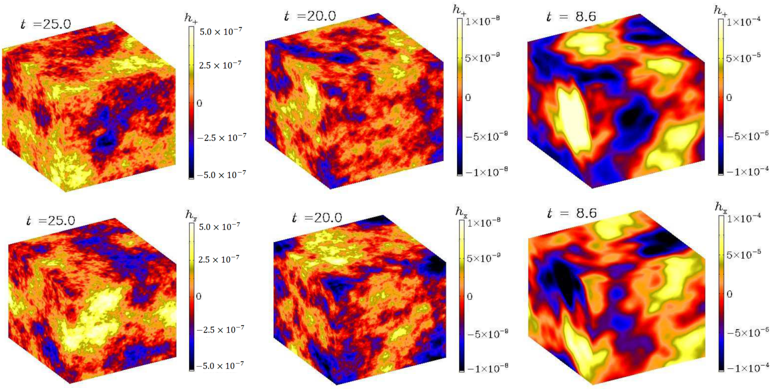
Evolution of magnetic strength and correlation length[28]



[28] A. Brandenburg *et al.* (incl. ARP), *Phys. Rev. D* **96**, 123528 (2017).

Numerical results for decaying MHD turbulence[29]

$1152^3, \Omega_{\rm M} \sim 10^{-2}$

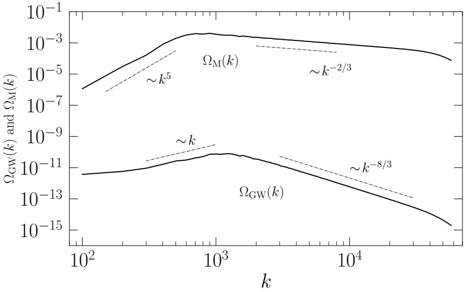


$k_{\rm M}/k_H = 300$ $k_{\rm M}/k_H = 60$ $k_{\rm M}/k_H = 2$

[29] ARP *et al.*, *Phys. Rev. D* **102**, 083512 (2020).

Numerical results for decaying MHD turbulence[30]

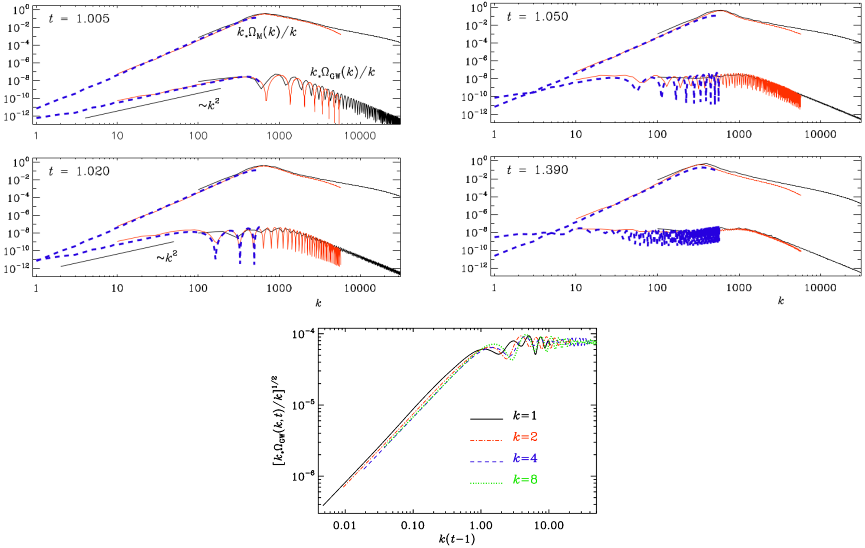
$1152^3, k_* = 2\pi \times 100, \Omega_{\rm M} \sim 10^{-2}, \sigma_{\rm M} = 1$



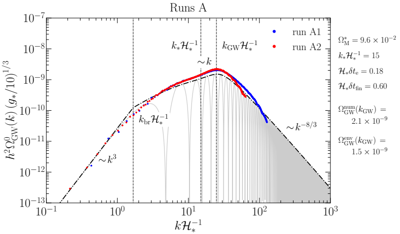
- **Characteristic k scaling in the subinertial range for the GW spectrum.**
- k^2 expected at scales $k < k_*$ and k^3 at $k < H_*$ according to the "top-hat" model (Caprini *et al.*, 2020).

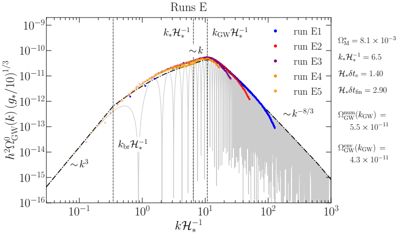
30 ARP *et al.*, *Phys. Rev. D* **102**, 083512 (2020).

Early time evolution of the GW spectrum



Numerical results for decaying MHD turbulence[31]





run	$\Omega_{\rm M}^*$	$k_*\mathcal{H}_*^{-1}$	$\mathcal{H}_*\delta t_{\rm e}$	$\mathcal{H}_*\delta t_{\rm fin}$	$\Omega_{\rm GW}^{\rm num}(k_{\rm GW})$	$[\Omega_{\rm GW}^{\rm env}/\Omega_{\rm GW}^{\rm num}](k_{\rm GW})$	n	$\mathcal{H}_*L$	$\mathcal{H}_*t_{\rm end}$	$\mathcal{H}_*\eta$
A1	9.6×10^{-2}	15	0.176	0.60	2.1×10^{-9}	1.357	768	6π	9	10^{-7}
A2	–	–	–	–	–	–	768	12π	9	10^{-6}
E1	8.1×10^{-3}	6.5	1.398	2.90	5.5×10^{-11}	1.184	512	4π	8	10^{-7}
E2	–	–	–	–	–	–	512	10π	18	10^{-7}
E3	–	–	–	–	–	–	512	20π	61	10^{-7}
E4	–	–	–	–	–	–	512	30π	114	10^{-7}
E5	–	–	–	–	–	–	512	60π	234	10^{-7}

31 ARP *et al.*, *Phys. Rev. D* **105**, 123502 (2022).

Analytical model for GWs from decaying turbulence

- Assumption: magnetic or velocity field evolution $\delta t_{\rm e} \sim 1/(u_* k_*)$ is slow compared to the GW dynamics ($\delta t_{\rm GW} \sim 1/k$) at all $k \gtrsim u_* k_*$.
- We can derive an analytical expression for nonhelical fields of the envelope of the oscillations[32] of $\Omega_{\rm GW}(k)$.

$$\Omega_{\rm GW}(k, t_{\rm fin}) \approx 3 \left(\frac{k}{k_*}\right)^3 {\Omega_{\rm M}^*}^2 \frac{\mathcal{C}(\alpha)}{\mathcal{A}^2(\alpha)}\, p_\Pi\left(\frac{k}{k_*}\right) \times \begin{cases} \ln^2[1 + \mathcal{H}_* \delta t_{\rm fin}] & \text{if } k\, \delta t_{\rm fin} < 1, \\ \ln^2[1 + (k/\mathcal{H}_*)^{-1}] & \text{if } k\, \delta t_{\rm fin} \geq 1. \end{cases}$$

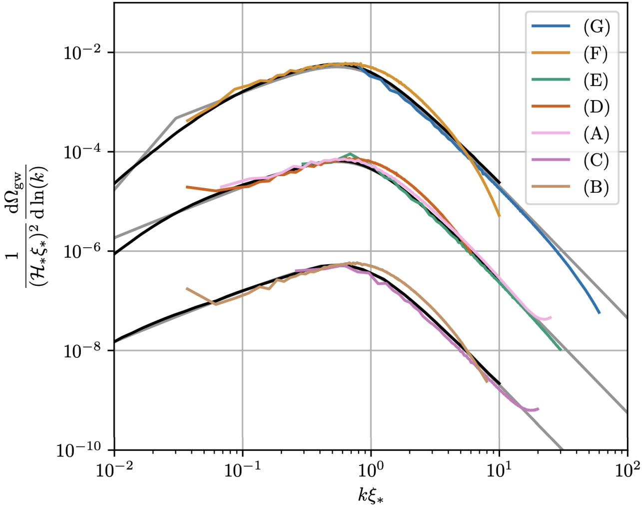
- p_Π is the anisotropic stress spectrum and depends on spectral shape, can be approximated for a von Kárman spectrum as[33]

$$p_\Pi(k/k_*) \simeq \left[1 + \left(\frac{k}{2.2 k_*}\right)^{2.15}\right]^{-11/(3\times 2.15)}$$

[32] ARP *et al.*, *Phys. Rev. D* **105**, 123502 (2022).

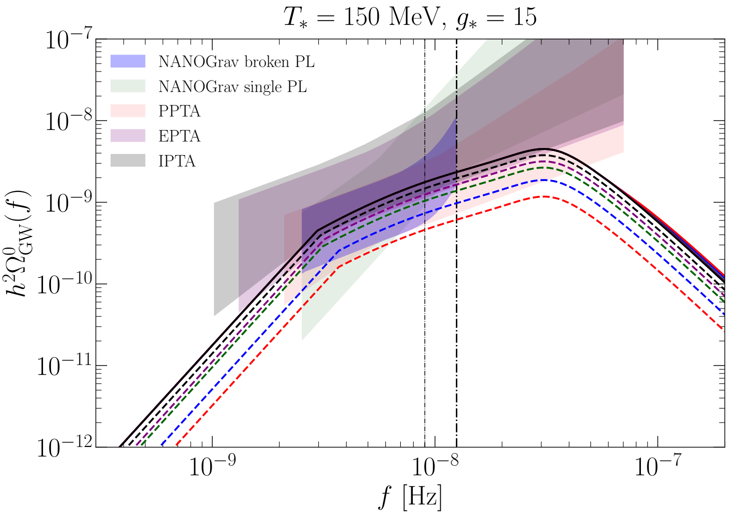
[33] ARP *et al.*, arXiv:2307.10744 (2023).

Numerical results for decaying HD vortical turbulence[34]



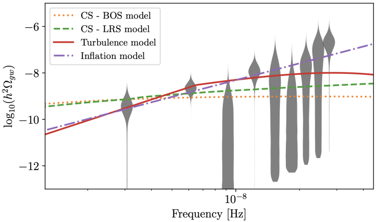
[34] P. Auclair *et al.*, *JCAP* **09** (2022), 029.

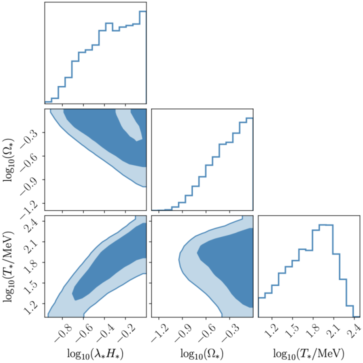
Using PTA to constrain primordial magnetic fields at the QCD scale[35]



35 ARP *et al.*, *Phys. Rev. D* **105**, 123502 (2022).

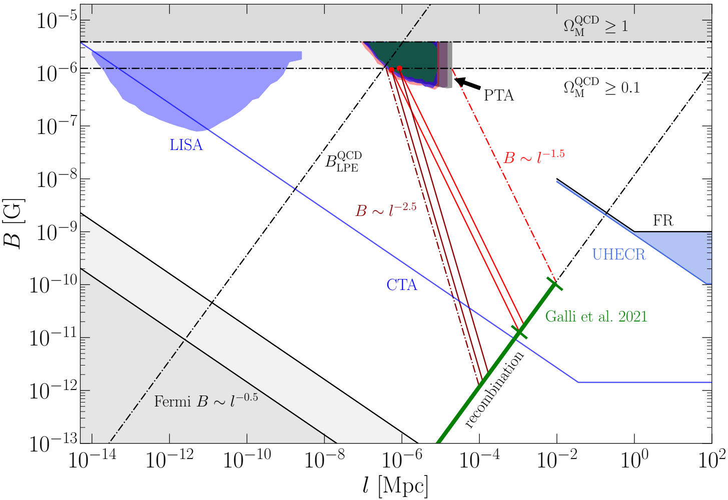
Primordial magnetic fields constraints with PTA[36]





36 [EPTA collab.] (incl. ARP), arXiv:2306.16227 (2023).

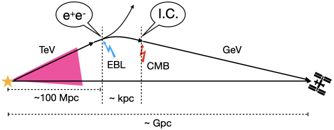
Primordial magnetic field constraints with PTA[37]



[37] ARP *et al.*, *Phys. Rev. D* **105**, 123502 (2022).

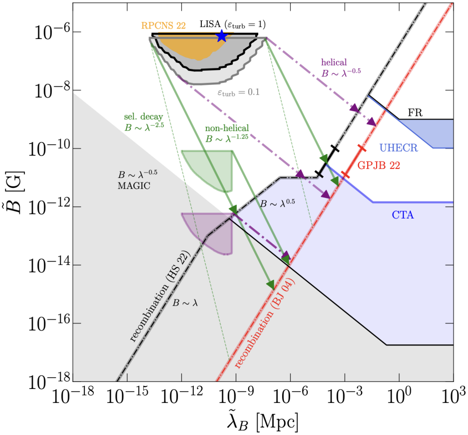
Multi-messenger studies of primordial magnetic fields[30]

- Primordial magnetic fields would evolve through the history of the universe up to the present time and could explain the lower bounds in cosmic voids derived by the Fermi collaboration.[31]



- Maximum amplitude of primordial magnetic fields is constrained by the big bang nucleosynthesis.[32]

- Additional constraints from CMB, Faraday Rotation, ultra-high energy cosmic rays (UHECR).



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[30] ARP *et al.*, arXiv:2307.10744 (2023).

[31] A. Neronov and I. Vovk, *Science* **328**, 73 (2010).

[32] V. F. Shvartsman, *Pisma Zh. Eksp. Teor. Fiz.* **9**, 315 (1969).

Conclusions

- Magnetic fields are ubiquitous in nature at all scales, from the smallest structures up to galaxies, cluster of galaxies, filaments, and potentially in cosmic voids of the LSS.
- Observations from γ-rays indicate the presence of intergalactic magnetic fields that could have their origin in the early Universe
- The existence of primordial magnetic fields at the epoch of recombination could potentially alleviate the Hubble tension.
- Magnetic fields can be produced in the early Universe during the period of inflation or during a cosmological phase transition (EW or QCD), leaving at the same time an imprint on the gravitational wave background (GWB).
- The large conductivity during the radiation-dominated era implies the development of MHD turbulence and therefore, the primordial plasma is non-linearly stirred by the magnetic field, modifying the dynamics with respect to a fluid dominated by acoustic motion.
- LISA, PTA, and next-generation ground-based detectors can be used to probe the origin of magnetic fields in the largest scales of our Universe, which is still an open question in cosmology.
- In summary, magnetic fields can be used as multi-messenger probes for early Universe physics combining astrophysical and cosmological observations, and can help us understanding high-energy physics, at scales inaccessible by other means.

Thank You!

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github.com/AlbertoRoper/cosmoGW

albertoroper.github.io